

## A catalogue of stellar 1612 MHz maser sources

P. te Lintel Hekkert, H. A. Versteeg-Hensel, H. J. Habing and M. Wiertz

Sterrewacht Leiden, P. O. Box 9513, 2300 RA Leiden, The Netherlands

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**Summary.** — We present a catalogue of stellar objects with 1612 MHz maser emission extracted from the literature published between 1963 and 1983. A total of 442 sources were found. Some unpublished data are also included. A nearly full reference list (updated to 1987) is given. We discuss briefly the nature of the sources.

**Key words :** OH/IR stars — masers — long-period variables.

### 1. Introduction.

In recent years very significant advances have been made in the understanding of the last stages of evolution : the evolution at the asymptotic giant branch (AGB).

Stars at the AGB lose mass at a large rate ( $\dot{M} > 10^{-5}$   $M_{\odot}$  yr $^{-1}$ ) through a circumstellar envelope (CSE) of dust and gas ; depending on whether or not the stellar envelope contains more oxygen than carbon the CSE is termed "oxygen-rich" or "carbon-rich". The chemistry of CSEs is very rich (e.g. Olofsson, 1985 and 1988) especially that of the carbon-rich stars. The most prominent molecules for the oxygen-rich CSE are H<sub>2</sub>O, OH and SiO. The conditions created by the constant and spherical mass loss are favorable for the existence of strong masers : over a large range in distance (typically 10<sup>16</sup> cm) the velocity field is very regular. OH masers can be observed throughout the Galaxy and even beyond : Wood *et al.* (1986) discovered the first extragalactic OH/IR star in the Large Magellanic Cloud. (For reviews of stellar masers see Reid and Moran, 1981 ; Herman and Habing, 1985 ; Sun and Kwok, 1987, Bedijn, 1987, 1988 and Alcock and Ross, 1986).

Rapid progress in the studies of CSEs occurred after the publication of the IRAS Point Source Catalogue (PSC ; IRAS team 1984). The various far infrared photometers made it possible to classify the AGB stars in a comprehensive way using two IRAS colours (S(60μm)/S(25μm) and S(25μm)/S(12μm)) e.g. Olmon *et al.* (1984), Zuckerman and Dyck (1987), van der Veen and Habing (1988). As a consequence more and more data have become available from follow up studies on PSC sources selected on their infrared colours – studies in various molecular lines : H<sub>2</sub>O (e.g. Engels *et al.*, 1984) ; CO (e.g. Arquilla *et al.*, 1986) ;

OH (Lewis *et al.*, 1985 ; te Lintel Hekkert *et al.* 1989 ; Eder *et al.*, 1988). Therefore this seems the right time to catalogue the "pre-IRAS" data and, ultimately, combine these with the information from the follow-up on the IRAS PSC. Most of the pre-IRAS data are scattered over a large number of individual publications. In an attempt to make a list (as complete as possible) of AGB stars we made a reference catalogue of stars with 1612 MHz masers. The 1612 MHz maser is the most observed and strongest radioline in circumstellar shells of AGB stars and the most common in oxygen-rich objects. The stellar nature of the source can easily be recognized from the shape of the 1612 MHz line profile. A further consideration for cataloguing this line is the increasing difficulty of the 1612 MHz observations, because of the disastrous interference from the GLONASS satellite system (see also : Carter 1986).

### 2. Description of the contents of the catalogue.

The first detections of stellar 1612 MHz masers were made by Wilson and Barrett (1968) in the direction of sources from 2.2 μm survey by Neugebauer and Leighton (1969) – the IRC. These sources were later identified with long period variables, especially with Miras and with supergiants. Around 1972/1973 the increasing quality of the receivers made it possible to make "unbiased" sky surveys without *a priori* information on potentially interesting sources. During this period, which lasted until IRAS flew (1983), the majority of the sources in our catalogue were discovered (about 350) (e.g. Caswell and Haynes, 1975 ; Johansson *et al.*, 1977 ; Bowers, 1978 ; Baud *et al.*, 1979). Most sources have no optical counterpart, but as infrared point sources they can be recovered (e.g. Schultz *et al.*, 1976 ; Evans and Beckwith, 1977 ; Fix and Mutel, 1984) ; hence the

Send offprint requests to : P. te Lintel Hekkert.

name OH/IR stars (Schultz *et al.*, 1976). 1612 MHz maser observations in the direction of Miras and (super-) giants (e.g. Bowers and Sinha, 1978) were repeated in the eighties, but only a small number of additional stars were discovered (e.g. Rieu *et al.*, 1979, Olnon *et al.*, 1980; Slootmaker *et al.*, 1985); these newly detected objects have weak 1612 MHz masers.

The catalogue is given in table I, the reference list in table II.

**2.1 SELECTION OF THE SOURCES FROM THE LITERATURE ; SELECTION CRITERIA.** — All the sources were taken from articles published in the years starting with 1963 up to and including 1983. A given 1612 MHz maser source was included when it was clear that the maser was not associated with or part of a larger (molecular) complex or HII region. Sources showing absorption in their 1612 MHz profile were excluded, because they are thought to originate in the expanding shells around compact HII (see Reid and Moran, 1981). We did not make any further selection on the basis of the 1612 MHz line profile shape or other available data on the source.

We included in the catalogue the unpublished sources found during the so-called "Strip Survey" by Olnon *et al.* (1981 and 1989). Only double peaked sources were considered by Olnon *et al.* (1981) and are quoted in the catalogue; the Olnon *et al.* (1989) paper will give the single peak sources as well.

**2.2 TYPES OF SOURCES.** — It is believed that almost all of these sources are CSEs surrounding long period oxygen rich variables losing mass at a high rate, but at a low outflow velocity (about  $15 \text{ km s}^{-1}$ ) and (averaged over several pulsation periods) in a spherically symmetric way.

A few groups of 1612 MHz point sources have uncertain nature. One group contains objects like W28 (A2) (Zijlstra and Pottasch, 1988; Harvey and Forveille, 1988) and K5-35 (Engels *et al.*, 1985), which may be HII regions. A second group contains sources thought to be either in the transition phase between AGB star and planetary nebulae (PN) or young PN. OH 349.36-0.20 (see Pottash *et al.*, 1987), VY 2 - 2 (see Seaquist and Davis, 1983) are two of the most prominent examples. A third group consists of 1612 MHz maser sources with very high outflow velocities; five such objects are now known (te Lintel Hekkert *et al.*, 1988); the best known (OH231.8 + 4.2) may also be an object in transition between AGB star and planetary nebula (Morris *et al.*, 1987).

It is clear that the more information comes available, the more complex consistent models will be needed to explain the observed phenomena; a particular example is Roberts 22 (see Allen *et al.*, 1980).

**2.3 THE 1612 MHz MASER INFORMATION : VELOCITIES AND PEAK FLUXES.** — As a rule 1612 MHz information has been taken from the discovery paper. We quote the peak fluxes and velocities of the two outermost maser spikes, the system velocity is assumed to be the average of the velocities of the two maser spikes. In the few cases that the source

exists of only one spike. The system velocity equals the velocity of the maser spike. Since the distribution of the expansion velocity is strongly peaked between 20 and 40  $\text{km s}^{-1}$  (Fig. 1), the error in the system velocity will, in the case of single peaks, be smaller than about  $20 \text{ km s}^{-1}$ , assuming that the second peak is too weak to be measured. The peak fluxes are given in Janskys. The epoch of the observations is quoted from the literature with the accuracy of one month.

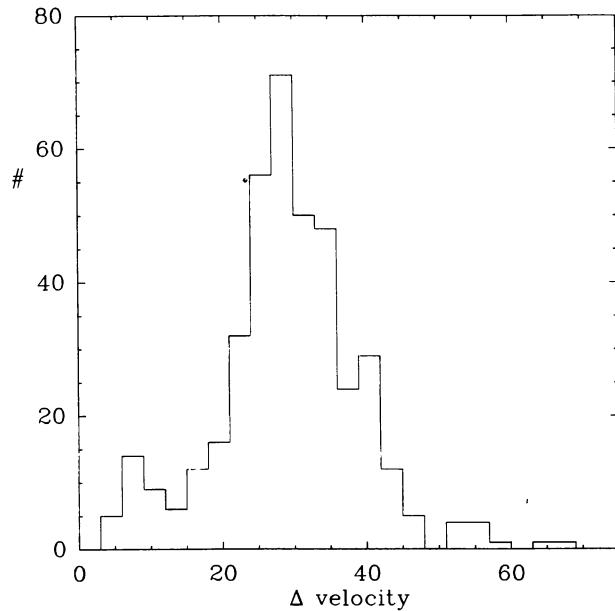


FIGURE 1. — Histogram of the expansion velocity of two peaked sources.

**2.4 POSITION INFORMATION.** — The most accurate position is quoted, which is not necessarily the position given by the discoverer. Warning: errors up to 10 arcminutes do occur in the quoted positions.

**2.5 IRAS IDENTIFICATIONS.** — We have searched in the IRAS catalogues for infrared counterparts on basis of positional information only and we give the IRAS source nearest to the catalogued position of the OH maser. A large number of the maser sources do not have accurate positions, so a more suitable tactic would have been to chose the nearest IRAS source with the "right" IR colours. We decided not to do so, since most of the maser sources are in, for IRAS, confused areas in the sky ( $|b| < 1^\circ 5$ ,  $|l| < 45^\circ 0$  and  $|b| < 2^\circ 5$ ,  $|l| < 10^\circ 0$ ) and it is not unlikely that the IR counterpart of the OH/IR stars was missed by the IRAS data point source processors. For the same reason we did not use the IRAS information in further determination of the nature of the source. Only 160 out of the 442 sources (36 %) of the sources in the catalogue could be identified with sources from the IRAS point source catalogue within  $0^\circ 3$ .

Some of the IRAS identifications could be checked by using the 1612 MHz OH maser surveys of the IRAS PSC by Eder *et al.* (1988), Lewis (private communications)

and te Lintel Hekkert *et al.* (1989). There are no misidentifications for sources that are within 0'.3 of an IRAS position.

### 3. References.

For every entry in the catalogue we tried to give a complete list of references, so the user will easily find other data known for his object. The references are updated to 1987, but are incomplete for literature published prior to the discovery of the source as a 1612 MHz maser source. The numbering of the references is not always sequential because the references given in table II are part of a larger, more general reference list on masers and AGB stars.

### 4. Reliability.

**4.1 THE DETECTION OF THE 1612 MHz MASER LINE.** — We included only 1612 MHz sources which were confirmed, either by the discoverers or by others. Since all AGB stars are variables and since they pump the 1612 MHz maser emission via the infrared continuum, the fluxes of the maser lines vary accordingly. The variation of the 1612 MHz maser flux can be as much as 1.5 magnitudes (Herman and Habing, 1985 ; OH127.9–0.0). In addition, there are a few stars known to have flares of maser emission, like U Orionis (Garrigue and Mennessier, 1980). The maser peaks associated with Mira variables and supergiants can vary by large factors (10 to 100), since the associated masers are not always saturated. Thus a source can be missed by one observer and found again a few years later by others.

The velocities of the maser peaks do not vary in velocity. During our research we used the velocities of the peaks as an extra identifier when comparing maser sources in different papers.

**4.2 THE STELLAR NATURE OF THE SOURCE.** — Four hundred ( $\approx 90\%$ ) of the sources show a two peaked 1612 MHz profile, characteristic of the spherical geometry of the CSE of AGB stars (Fig. 2a). The distribution of the ratio of the fluxes of the two peaks is very narrow and close to 1 (Fig. 3). A small percentage (about 1%) of these sources have very irregular line profile shapes (Fig. 2b). Some of these sources with irregular 1612 MHz line profiles could be objects with nonspherical mass outflow like VX Sgr (Chapman and Cohen, 1986).

The remaining  $\approx 9\%$  of the sources have only one maser peak detected. For most of these sources a possible second peak was missed due to poor-single-to-noise or strong baseline variations. There are a few well known “real” single peaks, like the young planetary nebula VY 2–2. For these objects it is thought that the second maser peak is absorbed by the ionised gas around the star.

A few sources lie possibly near or in HII regions. K 3-35 and W28(A2) are possible examples of (fast evolving) compact HII regions.

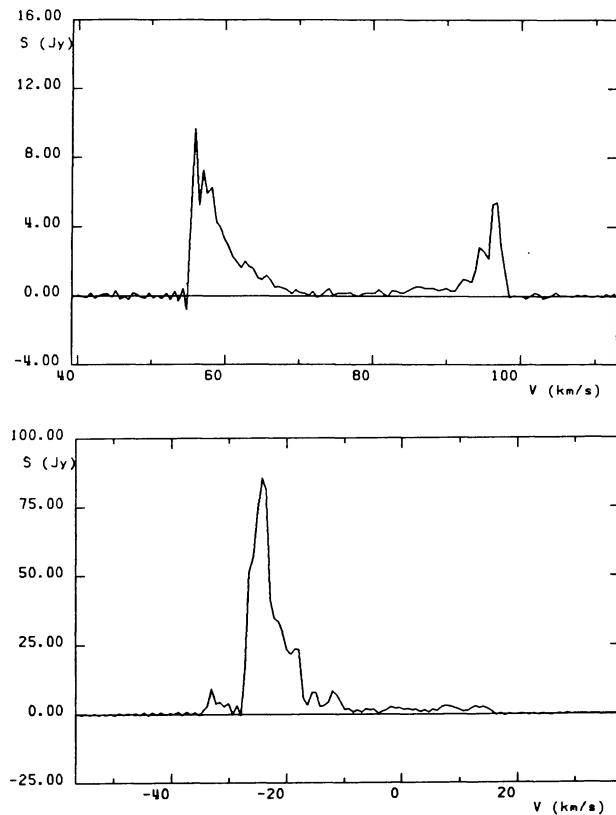


FIGURE 2. — a) A typical 1612 MHz maser line profile from an OH/IR star (# 362, OH 32.0–0.5; epoch August 1986); b) Roberts 22 (# 28), an example of a more complex 1612 MHz line profile (epoch July 1985).

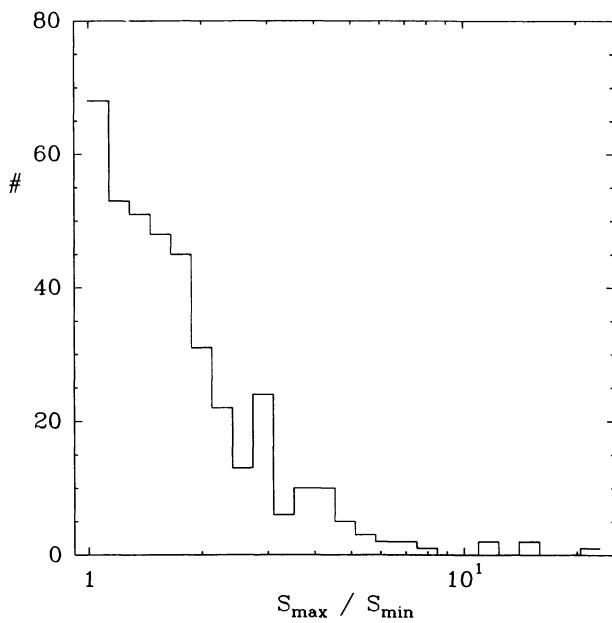


FIGURE 3. — Histogram of  $S_{\max} / S_{\min}$ .

**4.3 SKY COVERAGE AND 1612 MHz FLUX DISTRIBUTION.** — Figure 4 gives the spatial distribution of the catalogue sources. The figure shows clearly that not as much observing time has been devoted to the Southern part of the sky as the Northern. The sources at high galactic latitude are predominantly Mira variables and supergiants. Several

observers have searched (and found) OH/IR stars near the Galactic Centre (e.g. Habing *et al.*, 1983, Lindquist *et al.*, 1989). The radial velocity distribution is given in figure 5.

The distribution of the flux of the strongest maser spike is given in figure 6 ; it is clear that the sources in the catalogue form an inhomogeneous group.

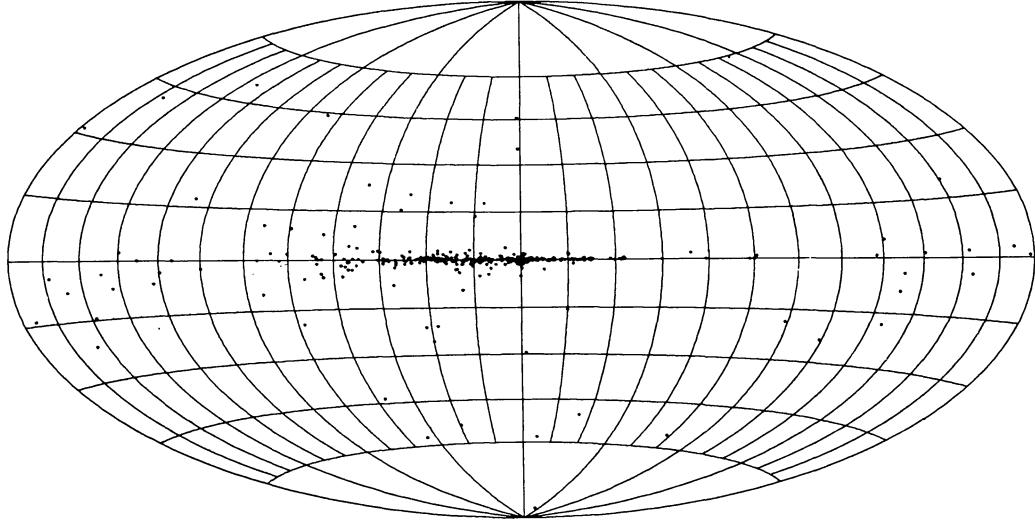


FIGURE 4. — The distribution of the sources in the catalogue in an Aitoff projection in Galactic coordinates.

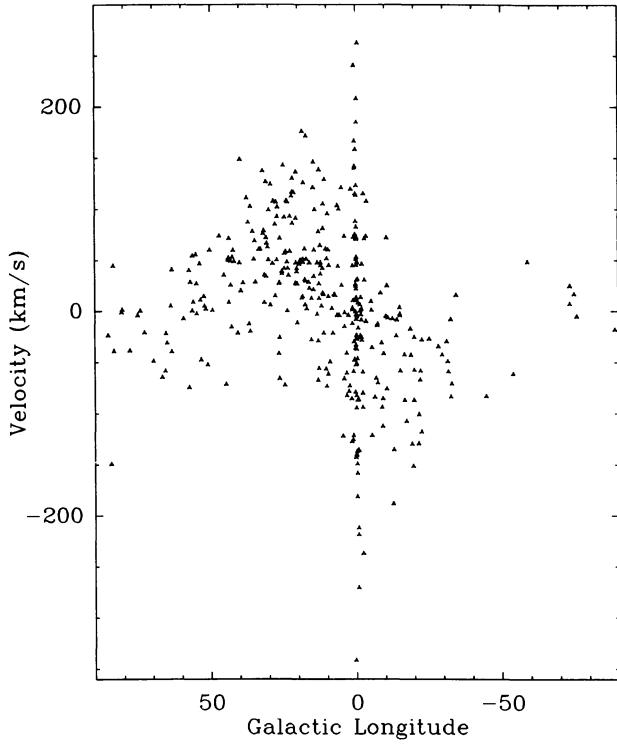


FIGURE 5. — The radial velocity distribution of the sources in the catalogue.

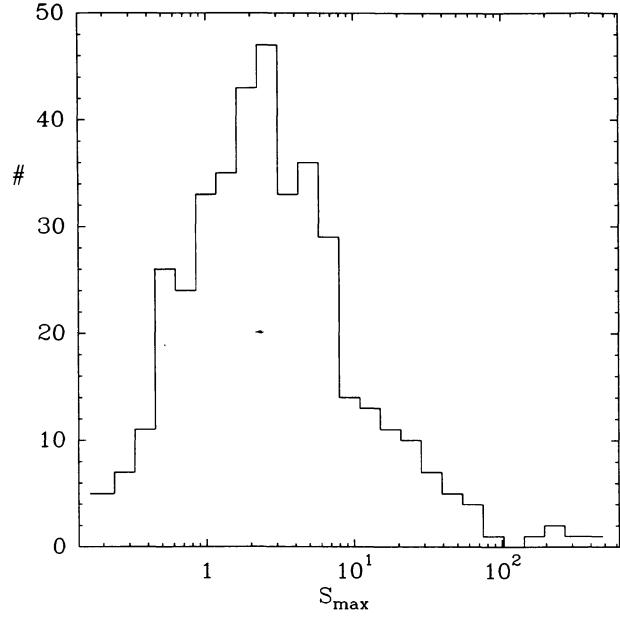


FIGURE 6. — Histogram of the 1612 MHz maser flux of the strongest 1612 MHz maser peak of the source.

## 5. Discussion.

**5.1 THE NATURE OF THE 1612 MHz MASER SOURCES.** — Assuming that the majority of the sources in the catalogue are AGB stars, it is interesting to see what stars on the AGB do and do not have 1612 MHz maser emission. (See for a recent review: Habing *et al.*, 1989). AGB stars include both oxygen-rich(e.g. Miras and OH/IR stars) and carbon-rich stars objects.

Mira variables are defined as optically identified, long period variables ( $P > 150$  d) (for a review on OH masers associated with Miras, see : Sivagnanam and Le Squeren 1988), those with a 1612 MHz maser are included in the catalogue. Most OH/IR stars are also long periods variables (but with longer periods than the visible Miras), but some have been observed not to vary at all or with a small irregular amplitude (the variation is measured by observing the 1612 MHz maser emission since the star is not visible due to the high opacity of the CSE).

Until recently there were no (visible) carbon stars known with a 1612 MHz maser. However, Willems and de Jong (1987) found nine objects in the IRAS LRS atlas identified with carbon star, that are either double stars of a very exceptional kind or objects in the transition phase between Mira variable and carbon star (Willems and de Jong, 1988 ; see also the discussion in Chan and Kwok, 1988). Two of these objects have been detected to have 1612 MHz maser emission (te Lintel Hekkert, private communications). The scenario given by Iben and Renzini (1983) for the transition from oxygen rich stars to carbon rich stars certainly leaves open the possibility that a few OH/IR stars (so defined because of the 1612 MHz maser emission) are in fact carbon stars with an oxygen rich CSE.

The S and symbiotic type of stars are the least understood classes of stars at the AGB ; no OH has been detected from

these stars (Norris 1985 *et al.*).

The classification of the supergiants is unclear ; some have 1612 MHz maser emission. The 1612 MHz profile of a supergiant is characterised by a large separation of the maser peaks ( $\Delta$  velocity  $> 40$  km s $^{-1}$ ), although there is a large overlap in expansion velocities with the OH/IR stars.

## 6. Conclusions.

We have produced a catalogue of 442 stellar maser sources at 1612 MHz. These are all the 1612 MHz maser emission sources discovered in the years 1963 up to and including 1983 (Tab. I). A nearly complete list of references (Tab. II) is given for the maser sources and this will enable the reader to easily find published data. The identification with IRAS PCS sources is only partly succesful, since most of the maser sources are in regions of the sky confused for IRAS.

## 7. Availability.

The catalogue is available on magnetic tape and can be requested from the authors. The tape contains the catalogue, the full reference list and access software.

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A number of people helped us in finding sources and removing mistakes from the earlier version of the catalogue. We thank Lauren Likkel for pointing out the problems with the IRAS identifications. Murray Lewis was the first to use the catalogue at the telescope and communicated (non-) identifications of IRAS sources with 1612 MHz maser sources and/or Mira variables prior to publication. We are especially indebted to Dr. J. H. Cahn, who lent us his unpublished OH maser catalogue. We thank Dr. L. L. E. Braes and Phil Maloney for a careful reading of the manuscript. Part of this research was financially supported through the Neherlands Organization for the Advancement of Pure Research (NWO) via grant 78-218.

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TABLE I

#	$\alpha$	$\delta$	1	b	V	$\Delta V$	S <sub>I</sub>	V <sub>I</sub>	S <sub>H</sub>	V <sub>H</sub>	Epoch	R <sub>es.</sub>	References	Name	IRAS name	Dist.
1	00:17:07.0	+65:42:54	119.725	3.317	-51.3	27.4	2.30	-65.0	2.10	-37.6	jan	84	0.15	8403 8601		001'17+6542 0.14
2	00:42:50.0	+68:54:36	122.442	6.317	-25.5	25.0	1.90	-38.0	1.60	-13.0	may	78	0.70	8003 8003 8407	IRC+70012	0042+6854 0.31
3	01:03:48.0	+12:19:51	128.641	-50.107	8.3	36.3	48.20	26.4	52.10	-9.9	nov	76	0.23	6801 6801 6901 7007 7015 7038 7111 7201 C1T3	0103+1219 0.02	
4	01:30:27.7	+62:11:30	127.815	-0.021	-55.0	22.0	32.00	-66.0	14.00	-44.0	jun	76	2.40	8602 7801 7425 7818 7502 7926 8101 8103 8105 OH127.8+0.0	01304+6211 0.13	
5	02:19:16.0	+58:21:30	134.623	-2.195	-38.5	29.0	1.10	-53.0	0.60	-24.0	oct	74	1.46	7511 7511 7417 7418 7519 7520 7535 7541 7638 S PER	02192+5821 0.20	
6	02:42:02.0	+12:06:22	161.473	-41.921	20.0	8.0	0.30	16.0	0.90	24.0	jan	76	1.80	7807 7807 7904 7926	RU ARI	02420+1206 0.07
7	03:20:40.7	+65:21:54	137.966	7.260	-37.5	19.0	8.00	-47.0	5.00	-28.0	jun	76	2.40	8105 7801 7818 7902 8009 8101 8103 8306 8402 OH138.0+7.2	03206+6521 0.40	
8	03:29:21.9	+60:10:25	141.718	3.522	-57.5	25.0	5.00	-70.0	1.00	-45.0	jun	76	2.40	8105 7801 7818 7902 8009 8101 8103 8306 8402 OH141.7+3.5	03293+6010 0.41	
9	03:50:43.4	+11:15:29	177.954	-31.413	34.2	32.9	2.20	17.7	3.30	50.6	nov	76	0.23	8590 7923 6801 6820 7007 7010 7015 7111 7202 NML TAU	03507+1115 0.02	
10	05:07:19.0	+52:48:53	156.437	7.834	2.9	33.7	4.30	-13.9	12.00	19.8	sep	70	0.70	7202 7202 7015 7111 7201 7207 7313 7331 7428 IRC+10050	05073+5248 0.11	
11	05:13:07.0	+45:30:48	162.952	4.338	-32.8	19.9	17.30	-42.8	1.20	-22.9	jan	85	0.58	7813 7324 7331 7418 7428 7520 7525 7601 7707 IRC+10050		
12	05:15:06.0	+63:12:51	148.282	14.563	51.5	29.0	15.00	37.0	11.00	66.0	may	78	0.70	8003 8003 8106 8311 8504 8528 8617	05131+4530 0.07	
13	05:32:21.0	-05:59:54	209.500	-19.750	7.5	0.31	7.5	nov	77	0.70	7139 8002	IRC+60154	05151+6312 0.20			
14	05:44:22.6	+27:07:09	181.754	-0.578	-8.0	0.40	-8.0	25.00	-39.0	aug	76	1.80	8590 7815 7904 7926 8003 8412	YY ORI	05443+2707 0.01	
15	05:52:50.0	+20:10:06	188.713	-2.495	-42.5	7.0	18.00	-46.0	75	0.16	7725 7725 7111 7127 7129 7131 7211 7313 7331 7331 Y ORI	05528+2010 0.20				
16	06:29:45.0	+40:45:08	174.114	14.122	-16.0	24.0	4.60	-28.0	2.20	-4.0	aug	69	0.58	7202 7202 7015 7111 7313 7717 7926 8103 8306 IRC+40156	06297+4045 0.16	
17	06:30:02.0	+60:58:54	154.309	21.520	-23.5	24.5	1.20	-35.7	0.42	-11.2	nov	77	1.12	7402 7903 7004 7732 7732 7926 8311 8504 IRC+60169	06300+6058 0.23	
18	06:31:59.0	+04:15:09	207.265	-1.808	12.9	0.26	12.9	nov	76	0.23	7923 7923 8207 8405	CRL961	06319+0415 0.01			
19	06:50:03.5	+08:29:02	205.576	4.126	-10.5	35.0	3.50	-28.0	5.50	7.0	feb	78	0.70	8590 8003 8105 8124 8617 8901	GX MON	06500+0829 0.01
20	07:05:28.0	-10:39:18	224.342	-1.287	52.0	26.8	0.41	38.6	0.19	65.4	feb	78	1.12	7903 7903 8111 8311 8438 8901	IRC+10143	07054-1039 0.16
21	07:20:54.0	-25:40:12	239.351	-5.068	22.4	56.8	200.00	-6.0	100.00	50.8	apr	77	1.16	7914 7914 6904 6911 6923 6924 6925 6927 6929 VY CMA	GL10174	07209-2540 0.37

TABLE I (*continued*).

#	$\alpha$	$\delta$	1	b	V	$\Delta V$	S <sub>I</sub>	V <sub>I</sub>	S <sub>H</sub>	V <sub>H</sub>	Epoch	Res.	References	Name	IRAS name	Dist.	
22	07:39:58.0	-14:35:43	231.833	4.217	15.9	23.0	0.15	4.4	0.23	27.4	dec 79	1.47	8004 8004 7102 7207 7309 7432 7433 7518 7528 OH0739-14	07399-1435	0.28		
23	07:44:34.0	-26:13:11	242.433	-0.693	82.7	9.5	1.90	77.9	1.00	87.4	mar 76	0.73	7729 7729 7904 7926 8111 8504	SS PUP	07445-2613	0.11	
24	07:58:30.0	-12:42:11	232.478	9.058	-15.6	16.3	0.57	-23.7	1.53	-7.4	jan 83	0.54	8590 8504 7202 7925	CRL1192 OH242-4-0.7 IRC-10184	07585-1242	0.01	
25	08:35:44.0	-10:13:40	235.322	18.097	17.8	43.3	1.40	-3.9	2.30	39.4	jul 76	0.73	7729 7729 7903 7926 8511 8605	U PUP CRL1274	08357-1013	0.13	
26	09:42:58.3	-21:48:04	255.803	23.356	39.0	24.0	7.60	27.0	5.70	51.0	sep 70	0.29	8590 7202 7015 7018 7111 7210 7313 7331 7447 IRC-20197	09429-2148	0.01		
27	10:18:54.9	-34:32:44	271.040	18.604	-17.9	6.8	0.16	-21.3	0.84	-14.5	jun 87	0.90	8306 8311 8412 8515 8617	8132 IW HYA			
28	10:19:45.1	-57:50:28	284.176	-0.788	-5.0	40.0	44.00	-25.0	3.00	15.0	aug 78	0.35	8013 8013 8126 8802 8901	7926 8023 8126 8802 8901	10189-3432	0.59	
29	10:28:43.0	-57:34:16	285.054	0.070	17.0	36.0	1.60	-1.0	2.20	35.0	mar 77	0.37	8108 8108 8203 8301 8404	7521 7611 7715	RCW 49 OH284-2-0.8	10287-5733	0.91
30	10:38:09.0	-58:17:48	286.501	0.059	25.0	32.0	2.10	9.0	2.30	41.0	mar 77	0.37	8108 8108 8203 8301	T VIR IRC-10264	10379-5817	1.50	
31	12:12:02.0	-05:45:30	286.545	55.659	7.3	7.7	0.38	3.5	0.36	11.2	feb 75	0.73	7904 7904 7926 8003 8412 8505 8901	12120-0545	0.09		
32	12:31:01.0	-62:33:43	300.930	-0.034	48.5	23.0	2.90	37.0	0.90	60.0	mar 77	0.37	8108 8108 8203 8301	T COM OH305-9-1.9	12310-6233	0.37	
33	12:56:12.0	+23:24:24	325.547	85.689	16.0	14.0	0.15	9.0	0.45	23.0	jan 76	0.80	7807 8077 7904 7926 8407	CRL1886	12562+2324	1.03	
34	13:15:58.9	-64:21:44	305.915	-1.915	-61.0	30.0	1.80	-76.0	1.30	-46.0	mar 76	0.89	8108 8108 8203 8301	13157-6421	1.55		
35	14:08:39.0	-07:30:42	334.719	50.122	-27.2	27.4	2.50	-40.9	1.90	-13.5	jul 76	0.73	7729 7729 7926 8054 8511	7904 7922 7926 7928 8012	14086-0730	0.09	
36	14:24:45.0	+04:53:53	352.664	57.971	-14.0	8.0	2.50	-18.0	7.50	-10.0	feb 78	0.73	8602 8003 7329 7417 7525 7619 7707 7713 7721 RS VIR	8070 7811 7817 7904 7926 8124 8126 8209 8311 IRC+C00243	14247+0454	0.36	
37	14:29:45.0	-60:10:53	315.217	0.013	-83.0	32.0	2.20	-99.0	1.90	-67.0	apr 77	0.37	8108 8108 8203 8301	8415 8610 8515 8901	14297-6010	0.59	
38	15:19:21.0	+31:32:45	49.472	57.175	0.8	6.5	4.00	-2.5	0.80	4.0	feb 78	0.70	8003 8003 7111 7127 7131 7211 7227 7313 7418 S CRB	7435 7447 7519 7520 7525 7713 7721 7811 7817 IRC+30272	15193+3132	0.12	
39	15:25:32.0	+19:44:13	29.515	53.477	6.0	14.8	2.50	-1.4	1.60	13.4	sep 70	0.29	8602 7202 6801 6901 7004 7015 7018 7110 7111 CIT7	7313 7324 7331 7453 7520 7525 7638 7707 7731 WX SER, OH327-1-0.3	15255+1944	0.09	
40	15:46:49.0	-54:20:10	327.099	-0.252	-71.0	34.0	7.80	-88.0	24.00	-54.0	jul 73	0.70	7507 7507 7706 7707 7818 8019	7732 7811 7817 7904 7922 7926 8003 8103 8124 IRC+20281	15468-5420	0.71	
41	15:47:39.0	-54:00:02	327.404	-0.066	-83.5	45.0	3.80	-106.0	2.80	-61.0	jul 73	0.73	7507 7507 7707 7707 7818 8019	7732 7811 7817 7904 7922 7926 8003 8103 8124	15476-5400	0.29	
42	15:50:16.0	-54:24:24	327.447	-0.622	-8.0	28.0	5.00	-22.0	19.00	6.0	jul 73	0.73	7507 7507 7707 7707 7818 7926 8118 8201 8301	7732 7811 7817 7904 7922 7926 8003 8103 8124	15502-5424	0.32	
43	15:51:33.0	-53:23:51	328.232	0.039	-59.0	40.0	104.00	-79.0	236.00	-39.0	jul 73	0.73	7507 7507 7707 7707 7818 7926 8118 8201 8301	8220 8301	15514-5323	0.82	
44	15:53:32.0	-53:29:01	328.405	-0.215	-49.0	38.0	4.90	-68.0	8.90	-30.0	jul 73	0.73	7507 7507 7707 7818 7926 8118 8201 8301	OH328-4-0.6	15535-5328	0.09	
45	15:55:18.0	-53:16:43	328.738	-0.227	-29.5	21.0	2.30	-40.0	5.00	-19.0	jul 73	0.73	7507 7507 7707 7818 7926 8118 8201 8301	OH328-7-0.2	15552-5316	0.33	
46	15:57:36.0	-12:12:35	358.436	29.525	-3.5	19.8	0.20	-13.4	0.20	6.4	feb 75	1.40	8463 904 7507 7507 7707 7818 7926 8118 8201 8301	FS LIB IRC-10329	15576-1212	0.35	
47	16:02:01.0	-51:57:54	330.368	0.100	-42.5	29.0	2.60	-57.0	1.80	-28.0	jul 73	0.73	7507 7507 7707 7818 7926 8118 8201 8301	OH330.4+0.1	16019-5157	0.37	
48	16:02:59.0	-30:41:33	344.993	15.725	-2.7	25.8	18.00	-15.6	10.50	10.2	jul 76	0.73	7729 7729 7926 8629	CRL1822	16029-3041	0.14	
49	16:09:40.0	-51:22:17	331.645	-0.257	-34.5	19.0	2.00	-44.0	8.90	-25.0	jul 73	0.73	7507 7507 7707 7818 7926 8118 8201 8301	OH331.6-0.3	16097-5122	0.34	
50	16:25:59.0	+34:54:36	56.370	43.533	54.3	24.5	3.20	-42.0	1.20	66.5	sep 69	0.58	7202 7202 7111 7133 7147 7148 7149 7601 7638 7713	IRC+30292	16260+3454	0.38	
51	16:33:26.0	-46:54:43	337.545	0.121	-118.0	52.0	7.10	-144.0	1.70	-92.0	jul 73	0.73	7507 7507 7707 7818 7926 8118	OH337.5+0.1	16333-4654	0.74	
52	16:33:47.0	-47:13:56	337.349	-0.140	-28.0	30.0	6.00	-43.0	9.50	-13.0	jul 73	0.73	7507 7507 7707 7818 8118	OH337.4-0.1	16337-4713	0.44	

N°3















TABLE I (*continued*).

#	$\alpha$	$\delta$	1	b	V	$\Delta V$	$S_I$	Sh	Vh	Epoch	Ref.	Name	IRAS name	Dist.		
356	18:46:07.2	-01:51:56	31.012	-0.220	126.7	29.8	7.00	111.8	9.30	141.6	aug 72	0.19	8502 7327 7512 7717 7818 8101 8105 8201 8203 OH31.0-0.2	18460-0151	0.38	
357	18:46:20.9	-01:43:30	31.164	-0.205	-28.0	10.10	-28.0				aug 72	0.19	8203 8301 8308 8415 8503 8506 8510	OH31.2-0.2	18464-0140	3.47
358	18:46:30.0	-01:20:33	31.521	-0.063	35.2	30.0	3.60	59.2	1.60	50.2	jul 74	0.19	8105 7717 7512 7818 8101 8201 8301	OH31.5-0.1	18465-0120	0.14
359	18:46:42.9	-02:38:14	30.396	-0.707	62.9	7.3	12.50	66.5	12.00	66.5	aug 72	0.19	8105 7327 7512 7630 7717 8309	OH30.4-0.7	18467-0238	0.24
360	18:46:48.0	-07:09:00	26.395	-2.799	-66.0	24.0	0.70	-78.0	1.10	-54.0	aug 77	1.21	7902 7302 8101	OH26.4-2.8	18469-0711	3.31
361	18:48:40.0	+01:53:24	34.640	0.944	28.5	29.0	1.40	2.70	43.0	sep 73	0.19	7717 7717 7818 7902 8101	OH34.7+0.9	18487+0152	1.19	
362	18:48:51.2	-01:07:29	31.985	-0.485	76.1	41.2	6.50	55.5	3.20	96.7	aug 72	1.86	8502 7327 7512 7717 7818 8101 8105 8301 8308 OH32.0-0.5	18488-0107	0.14	
363	18:49:25.0	-01:30:31	31.708	-0.787	78.8	25.7	3.80	66.0	6.60	91.7	jun 74	0.19	8415 8503 8506 8510 8511 8581	18494-0130	0.19	
364	18:49:35.0	+02:03:36	34.896	0.818	68.9	30.2	1.20	53.8	1.70	84.0	jul 74	0.19	8325 8412 8505 8603 8605 8901	OH34.9+0.8		
365	18:49:48.2	-00:17:52	32.828	-0.315	60.7	30.8	29.50	45.3	19.50	76.1	aug 72	0.19	8502 7327 7510 7512 7602 7621 7630 7634 7707 OH32.8-0.3	18498-0017	0.04	
366	18:49:52.0	+00:20:52	33.408	-0.032	60.5	31.0	1.00	45.0	3.00	76.0	jun 76	2.40	7801 7801 7818 8101 8510	OH33.4-0.0	18497+0022	1.88
367	18:51:30.0	+04:01:00	36.854	1.294	-12.5	15.0	2.70	-20.0	2.80	-5.0	aug 77	1.21	7902 7902 8101 8510	OH36.9+1.3	18512+0402	4.11
368	18:53:36.0	+07:24:00	40.101	2.384	47.5	37.0	1.20	29.0	2.00	66.0	aug 77	1.21	7902 7902 8101	OH40.1+2.4	18535+0726	2.76
369	18:54:09.0	+03:10:48	36.415	0.321	102.5	37.0	5.20	84.0	2.50	121.0	sep 73	0.19	7717 7717 7818 7902 8101 8506	OH36.4+0.3	18540+0302	0.99
370	18:54:56.0	+02:08:14	35.580	-0.333	78.0	28.0	18.00	64.0	14.50	92.0	oct 70	0.29	8602 7202 7021 7030 7111 7207 7313 7314 7405 OH35.6-0.3	18551+0159	9.02	
371	18:55:33.0	+01:44:24	35.298	-0.652	14.0		3.00	14.0			nov 72	0.19	7717 7717 7818 8101 8510 8629 OH1854+02	OH35.3-0.7	18556+0139	4.89
372	18:55:40.0	+01:35:57	35.187	-0.743	14.0	3.00	14.0				nov 72	0.19	8105 7717 8070 8309 8309 8309 OH39.7+1.5	18556+0136	0.58	
373	18:56:03.9	+06:38:49	39.713	1.495	20.0	32.0	20.00	4.0	34.00	36.0	jun 76	2.40	8602 7902 7729 7926 8009 8101 8103 8105 8511 8527	18560+0636	0.07	
374	18:56:32.0	+03:43:12	37.169	0.043	28.0		1.90	28.0			jun 74	0.19	8581 8503 8604 8605 8901	OH37.2+0.0	18567+0343	3.12
375	18:58:11.0	+06:16:48	39.630	0.859	19.8	32.5	1.30	3.5	3.40	36.0	sep 73	0.19	7717 7717 7818 8101 8510	OH39.6+0.9	18584+0616	4.30
376	18:59:36.2	+03:15:53	37.119	-0.847	87.1	28.8	10.00	72.7	15.00	101.5	sep 73	0.19	8502 7717 7512 7634 7818 7902 8009 8101 8105 OH37.1-0.8	18596+0315	0.02	
377	19:01:42.9	+06:08:44	39.916	0.018	148.6	29.2	8.00	134.0	11.90	163.2	sep 73	0.19	8602 7717 7512 7634 7818 7902 8009 8101 8103 OH39.9-0.0	19017+0608	0.18	
378	19:02:12.0	+00:46:00	35.201	-2.576	51.0	28.0	3.90	37.0	3.60	65.0	aug 77	1.21	7902 7902 8101	OH35.2-2.6	19020+0045	1.96
379	19:02:31.0	+03:33:00	37.708	-1.360	110.9	20.8	2.50	100.5	0.70	121.3	jul 74	0.19	7717 7717 7818 7902 8101 8506	OH37.7-1.4	19026+0336	4.04
380	19:03:57.7	+08:09:07	41.953	0.454	48.6	10.4	4.00	43.4	82.00	53.8	sep 70	0.29	7015 7202 7111 7207 7210 7331 7525 7724 7725 R AQL	19039+0809	0.07	
381	19:04:54.0	+10:13:00	43.889	1.207	50.0	32.0	1.00	34.0	0.70	66.0	aug 77	1.21	8103 8105 8115 8119 8124 8126 8306 8325 8415 8503 8505 8506 8510 8511 8516	19043+1009	8.70	
382	19:05:56.0	-22:19:12	14.664	-13.615	21.3	26.5	12.00	8.0	6.00	34.5	oct 70	0.29	7202 7702 7111 7313 7331 7447 7638 7707 7724 IRG-20540 V 2880 SGR	19059-2219	0.33	
383	19:06:34.5	+08:32:56	42.604	0.066	53.0	36.0	3.80	35.0	6.00	71.0	sep 73	0.19	8413 7717 7512 7818 7902 8101 8105 8201 8301 OH42.6+0.0	19065+0832	0.08	
384	19:06:43.8	+08:11:41	42.309	-0.133	59.5	32.7	5.00	43.1	24.00	75.8	sep 73	0.19	8602 7717 7512 7634 7818 7902 8003 8101 8105 OH42.3-0.1	19067+0811	0.28	
385	19:06:51.0	+08:40:55	42.754	0.068	76.0		6.00						OH42.75+0.07	19066+0838	3.92	
386	19:07:09.3	+09:47:02	43.764	0.514	8.6	23.7	3.70	-3.2	3.80	20.5	jun 74	0.19	8105 7717 7512 7818 7902 8101 8603	OH43.8+0.5	19071+0946	0.22
387	19:07:54.0	+09:00:00	43.156	-0.014	25.5	21.0	9.80	15.0	2.00	36.0	aug 77	1.21	7902 7902 6711 6803 6805 6809 6813 6814 W49	19078+0901	1.32	



TABLE I (*continued*).

#	$\alpha$	$\delta$	1	b	V	$\Delta V$	$S_I$	$V_I$	$S_h$	$V_h$	Epoch	Res.	References	Name	IRAS name	Dist.	
424	20:02:15.0	+28:53:50	66.762	-1.251	-64.5	25.0	1.00	-77.0	1.00	-52.0	jun	76	2.40	7801 7801 7818 7902 8101	OH66-8-1.3	20023+2855 1.83	
425	20:04:43.0	+12:48:10	53.360	-10.304	-47.5	7.0	1.40	-51.0	1.30	-44.0	may	78	0.70	8003 8003 7525 7638 7707 7713 7731 7904	7926 SY AQL	20047+1248 0.37	
426	20:07:46.0	-06:24:42	36.361	-20.406	-19.8	24.5	35.00	-32.0	3.00	-7.5	oct	70	0.29	7202 7202 7111 7313 7331 7447 7531 7638	7707 IRC-10529	20077-0625 0.58	
427	20:18:11.0	+22:34:12	63.436	-7.724	40.5	22.3	2.50	29.4	1.80	51.7	jan	84	0.15	7732 7926 8003 8132 8407 8515 8901	20181+2234 0.05		
428	20:25:55.0	-40:35:00	0.918	-35.176	-59.0	6.0	0.60	-62.0	0.50	-56.0	jul	75	1.40	7720 7720 7024 7106 7619 7707 7731 7904	7926 U MIC	20259+4035 0.29	
429	20:26:40.0	+38:57:01	77.916	0.223	-38.5	21.9	8.80	-49.5	5.70	-27.6	oct	70	0.58	7202 7202 6922 7007 7015 7018 7019 7111	7207 ON4	20266+3856 0.16	
														7313 7314 7324 7818 8101 8510	OH2026+38		
430	20:27:13.0	+35:35:40	75.268	-1.842	-4.0	24.0	6.00	-16.0	4.00	8.0	jun	76	2.40	8413 7801 7818 7902 8101	8510	OH77-9+0.2	
431	20:44:03.1	+04:12:36	51.116	-23.096	-52.4	4.8	0.90	-54.8	0.50	-50.0	mar	76	0.73	8590 7904 7202 7926	BR DEL	20440+0412 0.04	
432	20:44:33.0	+39:56:50	80.808	-1.909	-1.2	45.4	50.00	-23.9	170.00	21.5	jan	80	1.00	8463 8308 6801 6808 6820 6822	6901 6906 6907 NML CYG		
														6925 7001 7007 7009 7011 7014 7015 7016	7109 IRC+40448		
														7111 7113 7125 7134 7201 7209 7210 7223			
														7313 7321 7324 7403 7418 7428 7435 7442	7444		
														7447 7453 7519 7520 7528 7531 7533 7601	7624		
														7626 7724 7732 7817 7914 7936 7928 8101	8103		
														8107 8119 8124 8126 8128 8209 8212 8221	8306		
														8320 8406 8412 8416 8420 8427 8507 8510	8513		
														8615 8901	8604 8605 8615 8615 8618 8527 8581 8603	8604 8606	
433	20:49:10.3	+42:36:54	83.422	-0.889	-39.0	36.0	4.00	-57.0	1.00	-21.0	jun	76	2.40	8413 7801 7818 7902 8101	8510 8605 8629	OH83-4-0.9	
434	20:51:50.0	+44:46:00	85.380	0.126	-23.5	27.0	1.00	-37.0	2.00	-10.0	jun	76	2.40	7801 7801 7818 7902 8101	OH85-4+0.1		
435	20:53:00.0	+30:13:24	74.345	-9.407	0.5	25.0	1.50	-12.0	0.70	13.0	feb	78	0.73	8003 8003 7111 7202 7313 7519 7520 7525	7638 IRC+30464	20529+3013 0.10	
436	21:20:38.0	-40:55:00	1.328	-45.529	1.7	11.3	5.50	-4.0	2.70	7.3	oct	70	0.19	7106 7106 7215 7217 7331 7525 7707	7713 V MIC	21206-4054 0.03	
437	21:25:23.0	+36:29:00	83.657	-10.167	44.6	28.7	2.40	58.9	1.70	30.2	nov	76	0.23	7923 7923 7015 7111 7202 7313 7428 7638	7713 IRC+40483		
438	22:17:42.7	+59:36:16	104.908	2.414	-24.9	30.2	41.00	-40.0	39.00	-9.8	aug	77	0.67	8602 7902 7801 7818 7903 7926 8009	8101 8219 OH104.9+2.4	22177+5936 0.04	
														8310 8324 8402 8413 8428 8504 8510 8527	8603 AFGL2885		
														8604 8605 8901			
														8615 8901			
439	22:51:40.0	+08:37:54	80.572	-44.119	1.3	19.9	0.25	-8.6	0.10	11.3	nov	77	1.12	7903 7903 7202 7453 7923 7926 8311	8504 8901 IRC+10523	22516+0838 0.53	
440	22:55:35.4	+58:33:11	108.654	-0.860	-49.5	35.0	0.50	-67.0	0.50	-65.0	nov	77	0.70	8438 8438 8002 8504	8619 GL 2999	22556+5833 0.26	
441	23:41:41.0	+61:31:00	115.064	-0.044	-38.5	53.0	1.90	-65.0	2.30	-12.0	jan	72	1.00	7331 8419 7331 7519 7525 7541 7638	7814 7926 80417 8407 8528 PZ CAS	23416+6130 0.38	
442	23:42:34.0	+43:38:30	110.485	-17.342	-42.5	23.0	0.80	-31.0	0.55	-54.0	jan	82	1.12	8407 8407 8528 EY AND	IRC+40545 CIT14	23425+4338 0.31	

*Notations in table I :*

*Column 1 :* sequence number ;

*Column 2, 3 :* Right Ascension and Declination, equinox 1950 ;

*Column 4, 5 :* Galactic Longitude and Latitude ;

*Column 6 :* Stellar Velocity ( $\text{km s}^{-1}$ ) ;

*Column 7 :* velocity difference ( $\text{km s}^{-1}$ ) between the outermost maser spikes ;

*Column 8, 9 :* peak flux (Jy) and velocity ( $\text{km s}^{-1}$ ) of the red Doppler shifted velocity spike ;

*Column 10, 11 :* peak flux (Jy) and velocity ( $\text{km s}^{-1}$ ) of the blue Doppler shifted velocity spike ;

*Column 12 :* epoch of the OH measurement ;

*Column 13 :* velocity resolution ( $\text{km s}^{-1}$ ) of the OH measurement ;

*Column 14 to 22 :* references ; the position is quoted from the first reference ; the 1612 MHz maser information is taken from the second reference. Thereafter the references are given in chronological order ;

*Column 23 :* name (s) of the source ;

*Column 24 :* IRAS name of the nearest IRAS PSC ;

*Column 25 :* distance in arcminutes ('') between the IRAS position and the position given in column 2 to 5.

TABLE II. — *The references for table I. The references are given in chronological order. The sequence number is not complete, because table II is part of a more general list of references on masers and AGB stars. For convenience we also give the titles of the papers.*

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